

# A gravitational wave observatory operating beyond the quantum shot-noise limit

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**Around the globe several observatories are seeking the first direct detection of gravitational waves (GWs). These waves are predicted by Einstein's General Theory of Relativity (1) and are generated e.g. by black-hole binary systems (2). Current GW detectors are Michelson-type kilometer-scale laser interferometers measuring the distance changes between in vacuum suspended mirrors. The sensitivity of these detectors at frequencies above several hundred Hertz is limited by the vacuum (zero-point) fluctuations of the electromagnetic field. A quantum technology – the injection of squeezed light (3) – offers a solution to this problem. Here we demonstrate the squeezed-light enhancement of GEO 600, which will be the GW observatory operated by the LIGO Scientific Collaboration in its search for GWs for the next 3-4 years. GEO 600 now operates with its best ever sensitivity which proves the usefulness of quantum entanglement and the qualification of squeezed light as a key technology for future GW astronomy (4).**

Just as electromagnetic radiation is produced by accelerated charges, gravitational waves are generated by accelerated mass distributions, such as supernova explosions or neutron star and black hole binary systems spiraling into each other (2). GWs propagate at the speed of light and reveal themselves as an alternating stretching and compressing of space-time, transverse to their direction of propagation. The direct measurement of GWs is one of the greatest challenges of contemporary physics. Two neutron stars merging at the other side of our galaxy would produce a maximum space strain amplitude  $h$  here at the Earth of just  $h \approx 10^{-19}$ . The strain would be just  $h \approx 10^{-22}$  in the case that the same event is located somewhere in the Virgo cluster of galaxies being about 1000 times farther away. Michelson-type laser interferometers are suitable observatories to measure gravitational waves (5).

Currently a global network comprising 2 LIGO observatories in the USA (with an arm length of 4 km (6), not operational since Nov. 2010 due to upgrade activity), the Virgo project of the European Gravitational Observa-

tory (with 3 km arms (7)) and the German-British detector GEO 600 (with 600 m arms (8,9)) exists, with further observatories planned or proposed in Japan (10), Australia (11) and Europe (12). The targeted GW-frequency band extends from below 10 Hz to about a few kHz. An ideal probe for space-time disturbances requires test masses which are free-falling in the direction of the laser beams. Therefore, the interferometer mirrors are suspended as sophisticated multi-stage pendulums and are situated in ultra-high-vacuum systems.

To date no direct detections of gravitational waves have been made. However, upper limits on GW signal strength for certain classes of sources could be derived, see (13) and references there in. To realize gravitational wave astronomy with a daily event rate, a sensitivity improvement by at least one order of magnitude is required, see for example Ref. (12), thereby increasing the sensitive sky volume by 3 orders of magnitude. Science teams all over the world are currently addressing each instrumental noise source for future observatory generations. At lower and intermediate audio-band frequencies the sensitivity is, in practice, limited by seismic noise, suspension thermal noise and in future interferometers by photon recoil (below several tens of Hertz) and mirror thermal noise (in the range of hundreds of Hertz). At higher frequencies it is the quantum nature of light that inhibits a more precise measurement, because the counting statistics of the light particles themselves lead to a fluctuating interferometer output (shot-noise). This noise is caused by so-called “vacuum”, or “zero-point” fluctuations of the electro-magnetic field (3). The “classical” approach to improve the observatory's signal-to-shot-noise ratio is an increase of the circulating light power, since the signals produced by gravitational waves are proportional to the light power, whereas the shot-noise is proportional to only the square root of the power. However, a higher light power leads to a thermal deformation of the sensitive interferometer optics and an increasing radiation pressure noise level resulting in a practical upper limit for the optical light power applicable (12). Hence, additional technologies must be

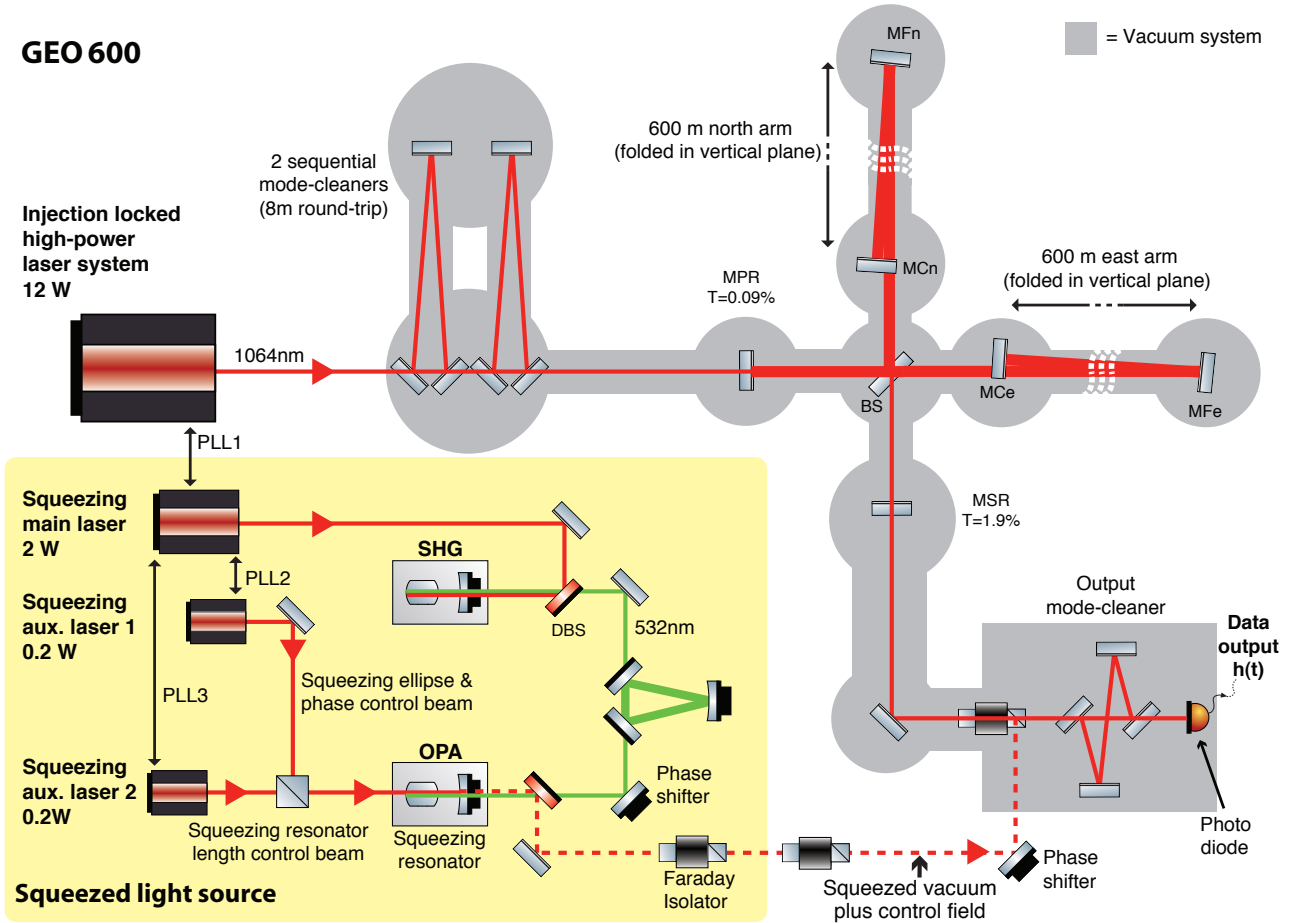


Figure 1: **Simplified optical layout of the squeezed-light enhanced GEO 600 observatory.** It consists of the conventional GEO 600 and the additional squeezed light source (yellow box, see method summary for details). The observatory has two singly folded arms with a total optical length of 2400 m. A GW passing from most directions will shorten one arm, while the length of the perpendicularly orientated arm is increased, and vice versa in the next half-cycle of a passing wave, producing a periodic power change of the output light that is detected by a photo diode. The observatory is operated such that almost all the light is back-reflected towards the 12 W input laser system, by keeping the interferometer output on a dark fringe by a control system. A power-recycling mirror (MPR) leads to a resonant enhancement of the circulating light power of 2.7 kW at the beamsplitter. Similar to the power-recycling technique, a partially transmissive signal-recycling mirror (MSR) is installed to further resonantly enhance the GW-induced signal at the interferometer's output. BS: 50/50 beamsplitter, SHG: second harmonic generator, OPA: optical parametric amplifier, DBS: dichroic beamsplitter, PLL: phase locking loop, MFe/MFn: far interferometer end mirrors (east/north), MCe/MCn : central interferometer mirrors, T: mirror transmissivity. All interferometer optics are suspended by multi-stage pendulums and situated in a vacuum system.

considered to push the sensitivity beyond this limitation (4).

Squeezed states of light (14) provide a way of increasing the sensitivity in the shot-noise limited region, *independently* of the circulating light power. Generally, a light field is described by two non-commuting physical quantities, the amplitude and phase quadratures. The minimum product of their uncertainties is limited by Heisenberg's uncertainty relation (HUR), which is also valid in the complete absence of photons, i.e. for a vacuum state. Vacuum states as well

as coherent states have noise equally distributed in the field quadratures. Only squeezed states – containing quantum correlated photons – show a noise below the vacuum noise level, however, due to the HUR, this is not possible for all quadratures of the state simultaneously. For a Michelson interferometer operated close to a dark output port, squeezed states can be utilized by injecting them into the observatory's signal output port and spatially overlapping them with the high power laser field at the beam splitter (3).

The squeezed quadrature has to be controlled such that, after being reflected off the interferometer, it is in phase with the readout (amplitude) quadrature of the observatory output light. This scheme produces path entanglement between the high-power light fields in the interferometer arms and reduces the photon counting noise on the photo diode in a way that can be explained only by photon correlations that are stronger than any classical correlation (4).

Since the first observation of squeezed light in 1985 (15), squeezed light sources have constantly been improved, recently reaching a factor of almost 13 dB below shot noise power (16). At frequencies in the GW detection band, the generation of squeezing remained an unsolved problem for a long time. Only recently, squeezing at Fourier-frequencies in the audio-band (17, 18) and a coherent phase control scheme for squeezed vacuum states could be demonstrated (19). In parallel, proof-of-principle experiments at higher frequencies have shown that small-scale Michelson-interferometer sensitivities can indeed be improved by squeezing (20, 21). Even though squeezed states are an ingredient for a multiplicity of quantum techniques like quantum teleportation (22), quantum memories (23) and many more, all of them are yet to mature from a proof-of-principle stage into a practical application. The increase of the GEO 600 sensitivity below its shot-noise limit by non-classical means is indeed the first practical application of this quantum technology, with the potential to become an integral part for all future generations of laser-interferometric gravitational wave observatories.

The German-British GEO 600 facility is one of the large-scale Michelson interferometers searching for gravitational wave signals. GEO 600 already uses a number of so-called *advanced techniques*, which are foreseen to be implemented in future upgrades of LIGO or Virgo (24). Fig. 1 shows a simplified layout of GEO 600. The first steps of the ongoing GEO high frequency (HF) upgrade program have already been included (25): GEO 600 is now operated in a tuned (to resonance with the laser carrier) signal-recycling mode with an optical homodyne detection scheme, also called *DC readout* (26, 27). Moreover, the carrier light transmitted by the SR mirror is filtered with an output mode cleaner cavity (OMC) (26), which suppresses technical modulation sidebands at radio-frequencies, as well as spurious light modes originating from mirror imperfections. Both filter effects assure that the detected beam is not contaminated with additional technical or quantum noise.

The lower left part of Figure 1 shows a simplified schematic of the squeezed light source which has been added in this work. The squeezed states of light are produced by parametric down-conversion inside an optical resonator, which contains periodically poled potassium titanyl phosphate (PPKTP) and which is pumped with single-mode continuous-wave light at 532 nm. The squeezed light beam

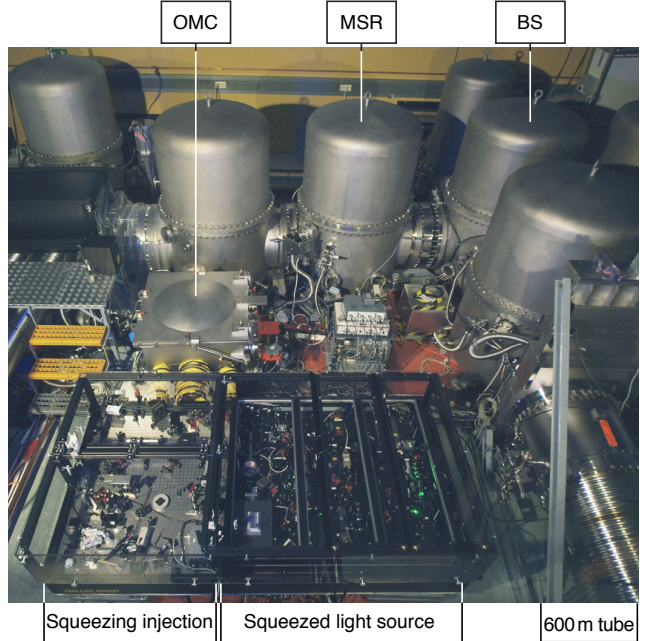
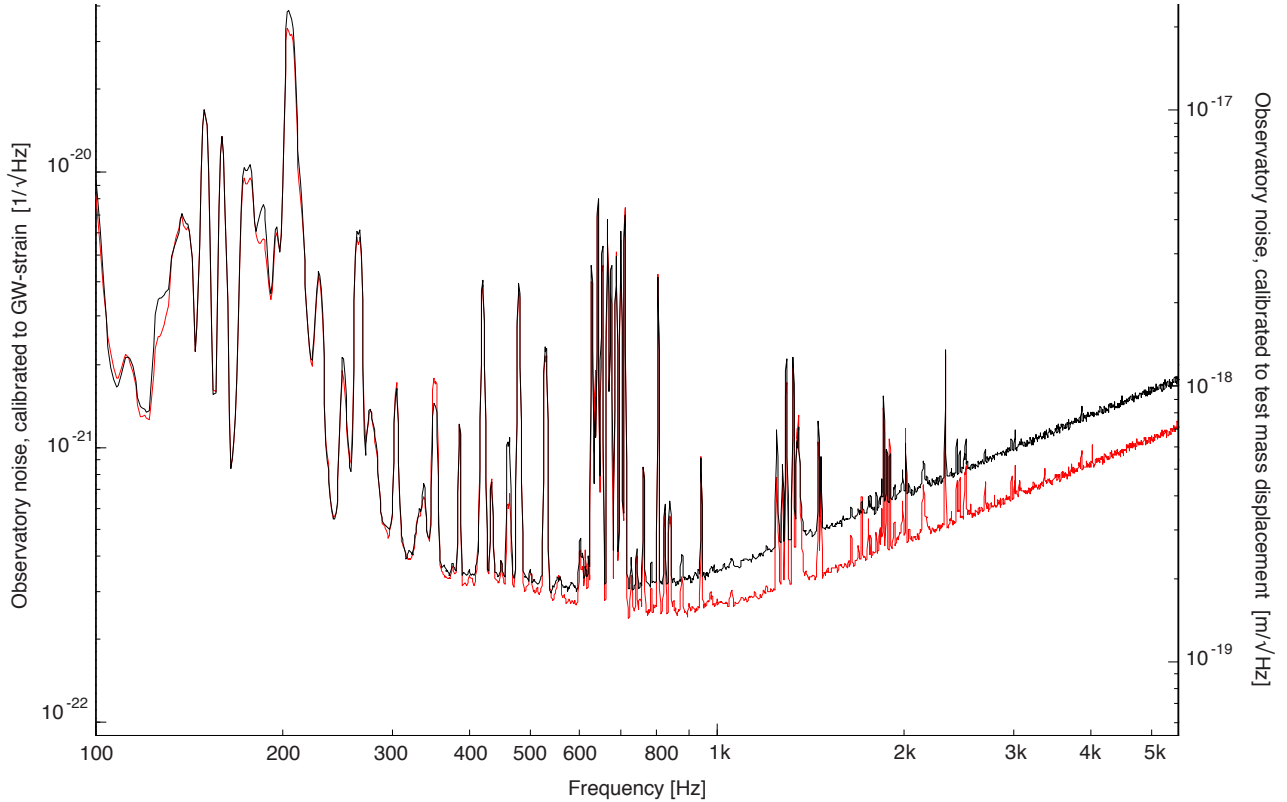


Figure 2: **View into the GEO 600 central building.** In the front, the squeezing bench containing the squeezed light source and the squeezing injection path is shown. The optical table is surrounded by several vacuum chambers containing suspended interferometer optics.

is phase locked to the 12 W GEO 600 laser and comprises squeezed vacuum states at frequencies from 10 Hz to above 10 kHz and a MHz control field for stabilization of the squeezed quadrature with respect to the GW signal (28). More details are given in the method summary. A view into the GEO 600 central building, showing the squeezed light source and the parts of the vacuum system, is presented in Figure 2.

Figure 3 presents the result achieved by this work: the quantum technology enhancement of an operating gravitational wave observatory. The injection of squeezed vacuum states into GEO 600 leads to a broadband noise reduction of up to 3.5 dB (black to red trace) in the shot-noise limited frequency band (above 700 Hz). The quantum noise at 3 kHz was reduced from  $1.0 \times 10^{-21} \text{ Hz}^{-1/2}$  down to  $6.7 \times 10^{-22} \text{ Hz}^{-1/2}$ . This corresponds to a factor  $1.5^3 \approx 3.4$  increase in detection rate for isotropically distributed GW sources in that frequency band. The squeezing enhancement has been successfully operated for several consecutive hours just limited by the current performance of the beam alignment. We see no principle limits for a quasi-continuous application of squeezed light, which is already planned for the next observational run of GEO 600. Due to the application of squeezed light the GW observatory GEO 600 has now achieved its best ever sensitivity since the implementation



**Figure 3: Nonclassical reduction of the GEO 600 instrumental noise using squeezed vacuum states of light.** The noise is calibrated to GW-strain and differential mirror displacement, respectively. In black the observatory noise spectral density is shown without the injection of squeezed light. At frequencies above 700 Hz GEO 600 is shot-noise limited; note that the slope in the kHz-regime is due to the normalization and the frequency dependent signal enhancement of GEO 600. An injection of squeezed vacuum states into the interferometer leads to a broadband noise reduction of up to 3.5 dB (red trace) in the shot-noise limited frequency band. The spectral features are caused by excited violin modes of the suspensions (600 - 700Hz and harmonics) as well as by calibration (160 Hz - 2.5 kHz) and OMC alignment control (250 - 550 Hz) lines. The broad un-resolved noise structures from about 120 Hz to 220 Hz are caused by insufficient seismic isolation of mirrors located between the signal recycling mirror and the output mode cleaner. Both traces shown were averaged over 4 minutes. The resolution bandwidth is 1 Hz for frequencies below 1 kHz, and 2 Hz at higher frequencies. Note that the noise reduction is independent of the averaging time used, such that the search for all kinds of potential GW sources (short- or long duration) benefits from the improvement.

of the advanced homodyne detection scheme. As expected, at Fourier frequencies below 700 Hz, squeezed light does neither reduce nor increase the current displacement noise level of about  $10^{-18} \text{ m}/\sqrt{\text{Hz}}$ . This observation makes us confident that a squeezed light improvement will extend to these frequencies as soon as the currently limiting technical noise will be reduced. Note, that quantum radiation pressure noise (29) is not expected to be significant at these frequencies at the present sensitivity.

The measured nonclassical quantum noise reduction in GEO 600 presented here is not limited by the squeezed light laser but by optical loss on the squeezed light during propagation in the interferometer. The 10 dB injected squeezed

state is degraded by photon absorption and scattering inside the GEO 600 signal recycling cavity and the output mode-cleaner, both contributing about 10 % loss. In addition, the non-perfect photo diode quantum-efficiency, the absorption of the Faraday isolators and auxiliary optics, and finally some residual mode mismatch cause an additional 20 % loss. All losses have been verified by independent measurements and provide an overall optical efficiency of  $\eta = 0.62$ . This leads to a reduction of the nonclassical noise suppression (in power) from  $V_{\text{sqs}} = 0.1$  (or 10 dB) to  $\eta V_{\text{sqs}} + (1 - \eta) = 0.44$  (or 3.5 dB), in excellent agreement with our results shown in Figure 3. Based on this, we are confident that future optical loss reductions will result

in a correspondingly higher squeezing factor. During the GEO-HF upgrade program in 2011 we expect a sensitivity improvement of up to 6 dB to be realized with squeezed light input. An even stronger impact through the application of squeezed light can be foreseen in future gravitational wave observatories, where state of the art optical technologies will allow for lower optical losses.

The results presented here show that squeezed light can improve operating gravitational wave observatories. We point out that squeezed light is also highly compatible with a future thermal noise reduction by means of cryogenic cooling of observatories (4), since in contrast to increasing the laser power, increasing the squeezing factor does not increase the thermal load on the mirrors. We therefore expect this innovative approach to become a key technology in making gravitational wave astronomy a reality, and we believe that squeezed light lasers in addition to high-power lasers are likely to be integrated into all future gravitational wave observatories.

## Method summary

Altogether four different laser frequencies are involved in the generation and coherent control of the squeezed vacuum states, see fig. 1. The main 2 W laser, which is phase locked to the 12 W GEO 600 laser, drives a second-harmonic generator (SHG). The green light from the SHG is filtered using a ring-resonator to attenuate high-frequency phase noise (30). The frequency up-converted field is subsequently injected into the squeezing resonator containing a nonlinear medium (periodically poled potassium titanyl phosphate, PPKTP) placed in a standing-wave half-monolithic cavity. Only 35 mW of the frequency doubled field are required to generate about 9 dB squeezing down to 10 Hz (28), via the process of parametric down-conversion (optical parametric amplification). In order to avoid any contamination of the squeezed light by laser noise in the audio band, two auxiliary lasers, frequency-shifted by several MHz, are employed for coherent control of the squeezed vacuum states (19, 28). Squeezing at Fourier-frequencies in the audio-band has been shown to be very sensitive to light back-scattered into the OPA (17–19). Therefore, the squeezed beam is guided through two Faraday isolator units before it is injected into the signal port of GEO 600.

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## Author Contributions

The activities of the LIGO Scientific Collaboration (LSC) cover modelling astrophysical sources of gravitational waves, setting sensitivity requirements of observatories, designing, building and running observatories and researching new techniques to increase the sensitivity of these observatories, and performing searches for astrophysical signals contained in the data. The principal investigators of the advancement reported here are H.G. and R.S. being responsible for GEO 600 and for the squeezed-light laser during the last 3 years, in which this experiment was prepared and conducted, respectively. In this period a great number of the LSC members contributed directly to the success of this project. H.V. and H.G. supervised the integration of the squeezed light laser into GEO 600. Together with A.K., they took and analyzed the data shown. The initial manuscript was written by a team involving those mentioned above together with R.S. The manuscript went into a two-stage LSC-wide review process, which was organized and led by R.F., T.R.C., M.H., and D.S.<sup>7</sup>. All authors approved the final version of the manuscript.

## Author Information

The authors declare no competing financial interests.